The Dark Energy Survey Supernova Program 5-year results, methodology, and future directions

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What is DES?



- Started observing in 2013
- DECam on the Victor M. Blanco Telescope
 - u**griz**Y filters
- Six years (758 nights)
 - 3 year preliminary SN results
 - 5 year final SN results
- Multi-probe survey
 - Supernovae
 - Gravitational Lensing
 - Galaxy Clusters
 - Baryon Acoustic Oscillations



Supernova Cosmology in Brief

Type Ia Supernovae (Observations)

 $\mu_{\rm SN}$ = m - M + $\Delta \mu$

What cosmological parameters make these equal?

Cosmological Prediction (Theory)

 $E(\hat{z}) = (\Omega_{R}(1+\hat{z})^{4} + \Omega_{M}(1+\hat{z})^{3} + \Omega_{K}(1+\hat{z})^{2} + \Omega_{\Lambda}(1+\hat{z})^{3(1+w)})^{0.5}$

 $D_{\rm L} = (1 + z) c / H_0 \int_0^z E^{-1}(z) dz$

 $\mu_{\rm Cosmo} = 5 \log_{10}(D_{\rm L} / 10[\rm pc])$

Calculated | Observable | Unknown

Tripp (1998)

Standardising Type la Lightcurves



Calculated | Observable | Unknown

Pippin - From lightcurves to cosmological parameter estimates

Designed by Sam Hinton Developed and maintained by me

Prepare & Simulate Data Prepare & simulate lightcurves for analysis

Lightcurve Fitting Fit <u>SALT</u> parameters to data and simulations using <u>SNANA</u>

Classification Photometrically classify lightcurves with <u>SuperNNova</u> or <u>Scone</u>

Contamination and Bias Correction Correct for contamination and bias using BEAMS with Bias Correction (<u>BBC</u>)

Systematic Covariance Matrix Compute a systematic covariance matrix, account for many systematic uncertainties

Cosmological Fitting Perform cosmological parameter inference with Bayesian methods like <u>COSMOSIS</u> or <u>WFIT</u> 3 year

Difference Imaging

SALT2

Spectroscopic

G10 Dust Model

CosmoMC

5 year

Scene Modelling Photometry

SALT3

Photometric

P21 Dust Model

Cosmosis

A Comparison of Results and Uncertainties

- Order of magnitude more supernovae
- Uncertainty ~1/2
- Systematics not yet dominant
 - SALT and dust majority

